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ABSTRACT

We present an H I study of the galaxy group LGG 351 using Widefield ASKAP L-band Legacy All-sky Blind Survey (WALLABY) early science data observed with the Australian Square Kilometre Array Pathfinder (ASKAP). LGG 351 resides behind the M 83 group at a velocity range \((cz)\) of \(\sim 3500–4800\) km s\(^{-1}\) within the rich Hydra-Centaurus overdensity region. We detect 40 sources with the discovery of a tidally interacting galaxy pair and two new HI sources that are not presented in previous optical catalogues. 23 out of 40 sources have new redshifts derived from the new HI data. This study is the largest WALLABY sub-sample to date and also allows us to further validate the performance of ASKAP and the data reduction pipeline ASKAPSOFT. Extended H I emission is seen in six galaxies indicating interaction within the group, although no HI debris is found. We also detect H I in a known ultra-faint dwarf galaxy (dw 1328–29), which demonstrates that it is not a satellite of the M 83 group as previously thought. In conjunction with multiwavelength data, we find that our galaxies follow the atomic gas fraction and baryonic Tully–Fisher scaling relations derived from the GALEX Arecibo SDSS Survey. In addition, majority of our galaxies fall within the star formation main sequence indicating inefficiency of gas removal processes in this loose galaxy group.

Key words: instrumentation: interferometer – galaxies: distances and redshifts – galaxies: groups: general – galaxies: interactions – galaxies: star formation.

1 INTRODUCTION

The environment in which galaxies are residing and evolving is thought to affect their fundamental properties, such as the morphology, colour, and star formation rate. Observations show that the fraction of late-type (spiral and irregular) galaxies in clusters increases with increasing redshift (Butcher & Oemler 1978; Oh et al. 2018a). Dressler (1980) highlighted the existence of a morphology–density relation, in which the fraction of early-type galaxies increases while the fraction of late-type galaxies decreases with increasing galaxy density. Studies have also revealed suppression of star formation in dense environments (e.g. Lewis et al. 2002; Kauffmann et al. 2004; Cortese et al. 2019). This evidence clearly shows that environmental effects play a significant role in galaxy evolution.

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As gas serves as a reservoir to fuel star formation, it is crucial to investigate the effectiveness of gas removal and accretion processes in environments of different density. To sustain ongoing star formation, a constant replenishment of gas is vital or else the galaxies would quench and become red and passive (Larson, Tinsley & Caldwell 1980). For gas accretion processes, cosmological models with cold and hot mode accretion have been postulated through numerical hydrodynamical simulations (Katz et al. 2003; Kereš et al. 2005). In the cold mode, fresh gas tunnels via filaments and accretes directly into the low- to intermediate-mass halo of galaxies. Hot mode accretion involves gas being shock-heated to high temperatures at the virial radius, and later cooling and falling into the galaxy centre. As a result, accretion processes via the cold mode are expected to dominate high-redshift (z > 1) galaxies, while the hot mode is expected to dominate high-density and high-mass (M > 10^{10.5} M_\odot) low-redshift (z < 1) galaxies (Dekel & Birnboim 2006; van de Voort et al. 2011).

The gas removal processes can occur through two main mechanisms: (i) ram-pressure stripping, when a galaxy passes through a dense intergalactic medium (IGM) and (ii) tidal stripping due to gravitational interaction. The first mechanism is well observed in high-density clusters (see e.g. Chung et al. 2009; Vollmer et al. 2012) but it has also been observed in low-density galaxy groups (Westmeier, Braun & Koribalski 2011; Rasmussen et al. 2012; Catinella et al. 2013; Brown et al. 2017). The second mechanism is more commonly observed in small groups, where relative velocities between galaxies are low (see e.g. Koribalski & López-Sánchez 2009; For, Staveley-Smith & McClure-Griffiths 2013; For et al. 2014). These mechanisms can result in a galaxy losing its hot halo and quenching of the star formation process. Quenching can occur in both cluster and group environments. Neutral hydrogen (H\textsc{i}) gas is a good tracer for studying these mechanisms in star-forming and H\textsc{i} rich galaxies as it is generally more extended than the stellar disc and is much more sensitive to any disturbance by the environment (Hibbard & van Gorkom 1996).

There have been studies of H\textsc{i} content and physical processes that affect galaxies in a wide range of environments, from voids to clusters and groups. Most notable studies are through single-dish H\textsc{i} surveys like the Northern hemisphere Arcetri Legacy Fast ALFA survey (ALFALFA; Giovannelli et al. 2005) and the H\textsc{i} Parkes All-Sky Survey (HIPASS; Barnes et al. 2001). The angular resolution of these surveys is limited to arcminutes which is usually insufficient to resolve gas discs and H\textsc{i} debris. Interferometers offer higher angular resolution but often at the cost of sensitivity to low column densities, and therefore require more observing time to reach the equivalent sensitivity of a single-dish telescope. Thus, they have been limited to small-scale target-specific surveys and resulted in biased samples for environmental studies. With new widefield H\textsc{i} surveys at sub-arcminute angular resolution commencing using radio telescopes such as the Australian Square Kilometre Array Pathfinder (ASKAP; Johnston et al. 2007), it will become possible to build an unbiased statistical sample of galaxies for investigating environmental effects.

The Widefield ASKAP L-Band Legacy All-sky Blind survey (WALLABY) is one of the key ASKAP science projects. The survey is designed to cover 3\pi steradian of sky (−90° < \delta < +90°) with an angular resolution of 30 arcsec and a spectral resolution of 4 km s\(^{-1}\) at z = 0. It is expected to detect the H\textsc{i} emission of more than 500,000 galaxies out to a redshift of 0.26 (Duffy et al. 2012). The detection rate will be an order of magnitude higher than that of the HIPASS and ALFALFA surveys. The estimated root-mean-square (RMS) noise level is 1.7 mJy per beam per 4 km s\(^{-1}\) (Koribalski et al., in preparation).

### 1.1 The Lyon Group of Galaxies 351

In this section, we provide some background information on the LGG 351 group, also known as NGC 5135 galaxy group. NGC 5135 has previously selected as part of the environmental study of galaxy groups that consist of Seyfert galaxies (Kollatschny & Fricke 1989, hereafter KF89). The confirmation of NGC 5135 group membership is based on spectroscopic observations, in which galaxies with radial velocities outside of ΔV \sim 1000 km s\(^{-1}\) are excluded. The KF89 catalogue for the NGC 5135 group consists of a total of seven galaxies. A revision of galaxy groups membership was carried out by Garcia et al. (1993) and Garcia (1993) (hereafter G93) using a combination of friends-of-friends (Huchra & Geller 1982) and Materne-Tully (Materne 1978; Tully 1980) methods. The joint method enables a more reliable way of building a galaxy group catalogue. G93 samples are selected from the Lyon-Meudon Extragalactic Database (Paturel et al. 1988) with cut-offs of B-band apparent magnitude of 14.0 and V\textsubscript{CMB} (recession velocity) ≤ 5500 km s\(^{-1}\). The NGC 5135 galaxy group was then re-grouped as LGG 351 in the G93 catalogue.

LGG 351 consists of 16 galaxies, with 6 galaxies overlapping with the KF89 identification. NGC 5126 is not listed in LGG 351 but a relatively small and nearby galaxy, PGC 46903, is included. Since the B-band magnitude of PGC 46903 is fainter than 14.0, it is likely a misidentification of NGC 5126 in G93. We also find that PGC 47574 is a duplicate of PGC 47573 for the given coordinates. In this paper, we consider NGC 5126 instead of PGC 46903 as (3) it is a loose group that spreads across 10 deg\(^{2}\) of sky, allowing the study of ‘pre-processing’ prior to cluster infall; (2) its stellar mass of ~10^{11} M_\odot is higher than previous WALLABY early science studies – NGC 7162 (Reynolds et al. 2019; hereafter R19) and NGC 7232 (Lee-Waddell et al. 2019; hereafter LW19) have group masses of ~10^{12} M_\odot – thus extending the range of WALLABY early science studies to higher density environments; (3) it is a loose group that spreads across ~10 deg\(^{2}\) of sky, covering a larger area and a higher redshift range than other WALLABY early science studies (Elagali et al. 2019; LW19; R19; Kleiner et al. 2019), and allowing us to confirm data quality over a wider frequency range.

This paper is structured as follows. Section 2 describes the ASKAP observations, data reduction and data quality. We summarize the source finding methodology and present the H\textsc{i} spectra, moment maps, and associated parameters in Section 3. In Section 4, we revisit group membership. In Section 5, we derive distances, stellar masses, H\textsc{i} masses, and rotational velocities. We compare the derived quantities with scaling relations in Section 6. In Section 7, we derive star formation rates and compare them with star forming main-sequence. Tilted ring fits to the velocity fields and an investigation of H\textsc{i} morphology and environment are presented in Sections 8 and 9. Further
The WALLABY early science phase 1 observations were carried out with two interleaving footprints (footprints A and B) with a square 6 × 6 beam pattern per footprint to achieve uniform RMS. Footprints A and B of the M 83 field are centred on J2000 coordinates $a_A = 13^\mathrm{h}37^\mathrm{m}54^\mathrm{s}$, $\delta_A = -29^\circ43'50''$, and $a_B = 13^\mathrm{h}39'58'53''$, $\delta_B = -30'10'46'36''$, respectively. We show the numbering of beams on the footprint in Fig. 1. In Fig. 2, we show the positions of the interleaved footprints with the shaded areas used in this study and the locations of LGG 351 galaxies, respectively.

The observations were carried out mostly during the day, and each observation began with a 2–3 h observation on the primary calibrator, PKS 1934–638. The total integration time on the M 83 field is 160.4 h with $\sim 135.8$ h being used for the data reduction. The observing log is given in Table 3. Each observation is given a scheduling block identification number (SBID).3

The M 83 field has also been observed as part of an extragalactic blind survey for very low column density H I (HIDEEP; Minchin et al. 2003). HIDEEP survey was carried out with single-dish 64 m Parkes telescope and covers a total area of $6' \times 10'$, reaching a column density sensitivity of $7.4 \times 10^{16}$ cm$^{-2}$ over a 20 km s$^{-1}$ velocity width in the central $4' \times 8'$ region. LGG 351 is located at the edge of the HIDEEP survey.

2.2 Data reduction

We reduced our data using ASKAPSOFT version 0.20.3. The details of the automated procedure will be published in Whiting et al. (in preparation). The customized procedure that we adopted is described in detailed in (Kleiner et al. 2019) so only a summary is given here. To reduce the processing resources required, we focused on a small range of channels, covering a bandwidth of 20 MHz. We also concatenated data from SBID 4619 and SBID 4620 into one data set.

### Table 1. Lyon Groups of Galaxies 351.

<table>
<thead>
<tr>
<th>PGC ID</th>
<th>Other name</th>
<th>RA (J2000) (h:min:s)</th>
<th>Dec. (J2000) (°:′:″)</th>
<th>mean $V_{hel}$a (km s$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>46585</td>
<td>ESO 444–G012</td>
<td>13:20:50.5</td>
<td>-29:28:49</td>
<td>4035</td>
</tr>
<tr>
<td>46902</td>
<td>NGC 5124</td>
<td>13:24:50.6</td>
<td>-30:18:25</td>
<td>4019</td>
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<tr>
<td>46974</td>
<td>NGC 5135</td>
<td>13:25:44.5</td>
<td>-29:49:59</td>
<td>3866</td>
</tr>
<tr>
<td>47078</td>
<td>IC 4248</td>
<td>13:26:47.7</td>
<td>-29:52:58</td>
<td>3902</td>
</tr>
<tr>
<td>47169</td>
<td>NGC 5150</td>
<td>13:27:36.6</td>
<td>-29:33:44</td>
<td>4117</td>
</tr>
<tr>
<td>47187</td>
<td>NGC 5152</td>
<td>13:27:50.7</td>
<td>-29:37:08</td>
<td>4227</td>
</tr>
<tr>
<td>47194</td>
<td>NGC 5153</td>
<td>13:27:54.7</td>
<td>-29:37:08</td>
<td>4068</td>
</tr>
<tr>
<td>47573</td>
<td>IC 4275</td>
<td>13:31:51.4</td>
<td>-29:44:01</td>
<td>4240</td>
</tr>
<tr>
<td>47489</td>
<td>NGC 5182</td>
<td>13:30:41.3</td>
<td>-28:09:03</td>
<td>4177</td>
</tr>
<tr>
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<td>13:22:23.6</td>
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<td>4276</td>
</tr>
<tr>
<td>47868</td>
<td>ESO 444–G021</td>
<td>13:23:30.3</td>
<td>-30:06:51</td>
<td>4084</td>
</tr>
<tr>
<td>46910b</td>
<td>NGC 5126</td>
<td>13:24:53.9</td>
<td>-30:19:48</td>
<td>4788</td>
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<tr>
<td>47145</td>
<td>IC 4251</td>
<td>13:27:24.5</td>
<td>-29:26:39</td>
<td>4180</td>
</tr>
<tr>
<td>47224</td>
<td>–</td>
<td>13:28:05.7</td>
<td>-29:25:29</td>
<td>3999</td>
</tr>
<tr>
<td>47574c</td>
<td>–</td>
<td>13:31:51.5</td>
<td>-29:44:02</td>
<td>4109</td>
</tr>
</tbody>
</table>

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<tr>
<th>PGCG ID</th>
<th>Other name</th>
<th>RA (J2000) (h:min:s)</th>
<th>Dec. (J2000) (°:′:″)</th>
<th>mean $V_{hel}$a (km s$^{-1}$)</th>
</tr>
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<td>47194</td>
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<td>-29:37:08</td>
<td>4068</td>
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<td>47573</td>
<td>IC 4275</td>
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<td>-29:44:01</td>
<td>4240</td>
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<tr>
<td>47489</td>
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<td>13:30:41.3</td>
<td>-28:09:03</td>
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<td>–</td>
<td>13:31:51.5</td>
<td>-29:44:02</td>
<td>4109</td>
</tr>
</tbody>
</table>

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1This is the default parameters for distances and cosmological corrected quantities in classic NASA/IPAC Extragalactic Database (NED) interface (Spergel et al. 2007).

2Documentation of ASKAPSOFT can be found at http://www.atnf.csiro.au/computing/software/askapsoft/sdp/docs/current/index.html

3Mean radial velocity derived from optical and/or H I spectra.

4Replacing PGC 46903 from the original G93 catalogue to PGC 46910 due to possible misidentification.

5Identified to be a duplicate of PGC 47573.
Table 2. ASKAP early science array configuration for the M 83/LGG 351 field.

<table>
<thead>
<tr>
<th>Observing period</th>
<th>Antennas</th>
<th>Minimum baseline (m)</th>
<th>Maximum baseline (m)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Dec 2016–Jan 2017</td>
<td>02,04,05,10,12,13,14,16,24,27,28,30</td>
<td>60</td>
<td>2300</td>
</tr>
<tr>
<td>Sept 2017</td>
<td>02,03,04,06,10,14,16,17,19,27,28,30</td>
<td>20</td>
<td>2300</td>
</tr>
<tr>
<td>Nov 2017–Jan 2018</td>
<td>01,02,03,04,05,06,10,12,14,16,17,19,24,27,28,30</td>
<td>20</td>
<td>2300</td>
</tr>
</tbody>
</table>

We ran ASKAPSOFT twice. In the first run, we processed all 36 beams with a 20 MHz bandwidth and stopped at the continuum stage. This yielded a 36-beam mosaicked continuum image per source SB. Each continuum image was used to assess the fluxes, spectral indices and positions of sources by comparing them to the NRAO VLA Sky Survey (NVSS; Condon et al. 1998) and the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al. 2003). This step was required because the early system has phase errors resulting in the phase centres not being exactly at the prescribed positions, which resulted in a small astrometric error of a few arcsec. The offsets for a given SBID were derived from the average offsets of the continuum sources from their true positions. Subsequently, we applied these offsets when we re-ran the full ASKAPSOFT the second time. We only reduced data from beams 4, 13–17, and 30–35 in the second run.

For each source SB, ASKAPSOFT first obtained the footprint information from the observational metadata and converted it to the corresponding beam centre position. Each source SB has an associated calibrator SB, which was used to perform the bandpass calibration. To perform the bandpass calibration, ASKAPSOFT split the calibration data into individual measurement sets (MSs) per beam for the selected channel range. All 36 beams were used for bandpass calibration. These calibration MSs went through autocorrelation, antenna and RFI flagging and were subsequently used to derive the bandpass solution for each beam (refer to Table 3 for flagged antennas).

The source data were also split into individual MS per beam for the selected beams and channel range. The bandpass solution was applied to the source MSs, calibrating the fluxes in each frequency channel by reference to the flux model of PKS B1934−638. Further gain calibration was done through self-calibration, and continuum images for each beam were then created with the calibrated, flagged and spectrally averaged MSs. Each continuum image was used as a model to carry out continuum subtraction in the UV-domain.

ASKAPSOFT then produced a spectral image cube for each beam. These ‘daily’ cubes, which could only be cleaned to a shallow level, were used to assess the data quality of each source SB. If an image cube was severely affected by artefacts, its calibrated source MS was excluded in the following customized imaging steps.

Final spectral-line imaging was performed separately using standalone scripts that used tasks in ASKAPSOFT. We first extracted a ~7 MHz sub-MS for each source MS and imaged each selected beam by combining useable sub-MS data (i.e. in the UV-domain) using the MSSPLIT and IMAGER tasks, respectively. The key parameters for the imaging included: 4 arcsec pixel$^{-1}$; Wiener robustness of 0.5; a Gaussian taper of 30 arcsec; and deconvolution with CLEAN using a single scale of one pixel (Högbom 1974). For cleaning, the major-cycle and minor-cycle CLEAN threshold was set to 3$\sigma$ and 4$\sigma$, respectively. One $\sigma$ corresponds to RMS/$\sqrt{N}$, where RMS is the typical single-night RMS of 7 mJy beam$^{-1}$ and $N$ is the number of combined nights. We created a (spectral-line) image cube per
Table 3. Observing and data reduction log.

<table>
<thead>
<tr>
<th>UT date (yyyy-mm-dd)</th>
<th>Calibrator SBID</th>
<th>Source SBID</th>
<th>Integration time (h)</th>
<th>Bandwidth (MHz)</th>
<th>Central frequency (MHz)</th>
<th>No. of antennas</th>
<th>Flagged antennas</th>
<th>Notes</th>
</tr>
</thead>
<tbody>
<tr>
<td>Footprint A</td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>2016-12-24</td>
<td>3017</td>
<td>3014</td>
<td>9.6</td>
<td>192</td>
<td>1344.5</td>
<td>12</td>
<td>16,30</td>
<td>–</td>
</tr>
<tr>
<td>2016-12-26</td>
<td>3026</td>
<td>3025</td>
<td>7.5</td>
<td>192</td>
<td>1344.5</td>
<td>12</td>
<td>13,16,30</td>
<td>Partially flagged on 13</td>
</tr>
<tr>
<td>2016-12-28</td>
<td>3034</td>
<td>3033</td>
<td>5.6</td>
<td>192</td>
<td>1344.5</td>
<td>12</td>
<td>16,30</td>
<td>–</td>
</tr>
<tr>
<td>2016-12-30</td>
<td>3040</td>
<td>3039</td>
<td>10.1</td>
<td>192</td>
<td>1344.5</td>
<td>12</td>
<td>16,24,30</td>
<td>Partially flagged on 24</td>
</tr>
<tr>
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<td>3048</td>
<td>3047</td>
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<td>13,16,30</td>
<td>Partially flagged on 13</td>
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<tr>
<td>2017-09-23</td>
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<td>4371</td>
<td>9.0</td>
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<td>1368.5</td>
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<td></td>
</tr>
<tr>
<td>2017-09-27</td>
<td>4396</td>
<td>4387</td>
<td>1.9</td>
<td>192</td>
<td>1368.5</td>
<td>12</td>
<td>–</td>
<td></td>
</tr>
<tr>
<td>2017-09-28</td>
<td>4401</td>
<td>4400</td>
<td>9.0</td>
<td>240</td>
<td>1320.5</td>
<td>16</td>
<td>14</td>
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<td>Footprint B</td>
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<td></td>
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<tr>
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<td>3021</td>
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<td>12</td>
<td>16,30</td>
<td>–</td>
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<tr>
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<td>3029</td>
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<td>16,30</td>
<td>–</td>
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<tr>
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<td>16,24,30</td>
<td>Partially flagged on 24</td>
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<tr>
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<td>3044</td>
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<td>–</td>
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<tr>
<td>2017-09-26</td>
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<td>12</td>
<td>–</td>
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<td>16</td>
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<tr>
<td>2017-12-16</td>
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<td>1320.5</td>
<td>16</td>
<td>–</td>
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<tr>
<td>2018-01-04</td>
<td>4936</td>
<td>4935</td>
<td>8.0</td>
<td>240</td>
<td>1320.5</td>
<td>16</td>
<td>–</td>
<td></td>
</tr>
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</table>

Beam. Subsequently, the residual continuum emission in the cube was removed by fitting a low-order polynomial to each spectrum. This yielded 12 image cubes per footprint and 24 image cubes in total. Finally, we mosaicked image cubes using the LINMOS task, where the primary beam correction was also applied. The image cube of beam 31 (footprint B) was excluded from the mosaic due to severe artefacts caused by solar interference (see Section 2.3). Excluding short baselines to mitigate against solar interference was avoided so as not to miss out on extended emission.

The final mosaicked cube covers a frequency range in the barycentric frame of 1396.57–1403.56 MHz (corresponding to $V_{\text{opt}} = 3599.1–5115.8$ km s$^{-1}$) and has a synthesized beam of 45 arcsec $\times$ 35 arcsec.

2.3 Data quality assessment

Statistical analysis was performed on the combined image cubes to assess the data quality of the ASKAP-12 data. RMS flux density values per channel per beam were measured. We find average RMS values of 2.2 and 2.6 mJy for beams in footprint A and B, respectively. These values are slightly better than the theoretical RMS of $\sim 2.8$ mJy for the given on-source integration time (assuming $T_{\text{sys}}/\eta$ of 75 K).

To characterize the distribution of noise in our image cubes, we examined the distribution of flux values for all voxels in each cube and calculated 1 and 99 percentile values of the distribution. The distribution of the noise is well represented by a Gaussian. The noise level can also be characterized by the 1 percentile rank of the pixel flux distribution to avoid bias by the contribution from the H1 sources (see e.g. Wong et al. 2006). In Fig. 3, we show the map of 1 percentile noise levels for beams in footprints A and B. Higher negative values indicate excessive negative flux, which is due to issues with bandpass calibration, sensitivity drop off towards corner beams and/or negative sidelobes. The text in each rectangle represents beam number.

Observations of the field were planned to avoid certain angles from the Sun in order to minimize solar interference. However, image stripes, which are a signature of solar interference, are seen in some of our image cubes. We find that they are better characterized by the maximum flux density. In Fig. 4, we show the maximum flux density as a function of frequency for beams 14 and 31 in footprint A and beam 31 in footprint B. The RMS values do not show such difference or variation with frequency, but the maximum

Figure 3. 1 percentile rank noise levels for combined beam image cubes of both footprints. Darker, negative values indicate issues with bandpass calibration, sensitivity drop off for corner beams and/or negative sidelobes. The text in each rectangle represents beam number.
Figure 4. Maximum flux density (Jy) as a function of frequency for the beam 14 and 31 image cubes in footprint A and the beam 31 data cube in footprint B. There is a significant variation in max flux density caused by the solar interference.

flux density in beam 31 for footprint B varies significantly by $\sim 0.4$ Jy across the whole frequency range. Using the maximum flux density as a metric, we identify the SBID 4623 data set as the main contributor to the stripes. This observation was carried out closest to the Sun, with beam 31 located at a Sun angle of 40°.

3 SOURCE FINDING AND CATALOGING

We search for HI sources in the mosaicked cube using the Source Finding Application (SOFIA; Serra et al. 2015). SOFIA is developed with ASKAP and future HI surveys in mind. WALLABY early science data are well suited to test SOFIA and to aid improving its algorithms. For the input parameters, we set a sub-cube region to exclude the first and last 10 channels due to low signal to noise (S/N) in these channels that do not contain all the combined data due to Doppler corrections. We use the local RMS for the noise scaling and a $5\sigma$ threshold for the smooth + clip source detection algorithm. We merge detected voxels into objects with a merging radius of 8 pixels and five spectral channels, and apply mask dilation for all detected sources to recover their total HI flux.

SOFIA generates outputs of individual cubelets and associated mask cubes for detected sources. For each source, a spectrum, integrated HI intensity (0th moment), velocity field (1st moment), velocity dispersion (2nd moment), and position–frequency maps are also generated based on the mask cube. To verify the detected sources, we follow these steps:

(i) Overplot the source positions on to images of the same region from the NVSS and the SUMSS to identify incomplete continuum subtraction. It was a known issue with this version of ASKAPSOFT that bright continuum sources ($\gtrsim 0.5$ Jy) are unlikely to be subtracted completely.

(ii) Cross match with the NED and SIMBAD databases (Wenger et al. 2000). The search radius is first set to 45 arcsec and then 60 arcsec if no source is found within 45 arcsec (~1 synthesized beam size).

(iii) Check the remaining detections by eye using the corresponding cubelet, 0th and 1st moment maps. A genuine detection generally has a coherent structure and velocity field. If selected, we further cross check them with optical images.

The final catalogue consists of 40 HI sources, which includes a newly identified tidally interacting galaxy pair (WALLABY J133002−272832) and two sources (WALLABY J133201−294119 and WALLABY J133237−295743) that do not have optical identifications. There are 15 galaxies in LGG 351 with five outside the footprint coverage area. Seven of the remaining 10 galaxies have HI detections. The on-sky distribution of 40 HI sources with their numbering is shown in Fig. 5. There are a few partial HI detections (e.g. only half of the double-horn profile is detected), which is due to the detection limit of the survey and/or higher noise at the edge of the beam where the source is located. We will exclude the tidal interacting galaxy pair and partial HI detections from the following statistical analysis, leaving a tally of 36 HI detections. There are only 5 and 16 of them detected in HIPASS and HIDEEP, respectively. The properties and parameters of each source are given in Tables 4 and 5. The derivation of parameters is described in the following sections.

3.1 HI spectra and integrated flux

SOFIA extracts the spectrum of each source by integrating the flux densities of all spatial pixels in each spectral channel. We plot the spectra in the optical velocity frame, $V_{\text{opt}} = cz$, where $c$ is the speed of light and $z$ is the redshift. The integrated flux ($S_{\text{int}}$) is then the integral of the flux densities of all channels in the spectrum. We

4Available at https://github.com/SoFiA-Admin/SoFiA

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### Table 4. Source catalogue and derived parameters.

<table>
<thead>
<tr>
<th>ID</th>
<th>Designation</th>
<th>Other ID</th>
<th>α (J2000) (°)</th>
<th>δ (J2000) (°)</th>
<th>$v_{lsr}$ (MHz)</th>
<th>$V_{rot}$ (km s$^{-1}$)</th>
<th>$V_{CMB}$ (km s$^{-1}$)</th>
<th>$D_H$ (Mpc)</th>
<th>$D_L$ (Mpc)</th>
<th>$N_{HI,sen}$ ($10^{19}$ cm$^{-2}$)</th>
<th>$S_{int}$ (Jy Hz)</th>
<th>$\log M_{HI}/M_\odot$</th>
<th>Note</th>
</tr>
</thead>
<tbody>
<tr>
<td>1</td>
<td>WALLABY J132514−303304*</td>
<td>ESO 444−30</td>
<td>201.30863</td>
<td>−30.55121</td>
<td>4.044</td>
<td>4309</td>
<td>4597</td>
<td>63.0</td>
<td>63.9</td>
<td>7.0</td>
<td>7.0</td>
<td>3700.2</td>
<td>&gt;8.89</td>
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<tr>
<td>2</td>
<td>WALLABY J132533−303521</td>
<td>LEDA 721502</td>
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<td>−30.08933</td>
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<td>4409</td>
<td>4697</td>
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<tr>
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<td>FLASH J132644.53−300029</td>
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<td>−30.00881</td>
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<td>4324</td>
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<td>3.7</td>
<td>3.7</td>
<td>3029.8</td>
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</table>

Note: This table is available in its entirety as Supporting Information with the electronic version of the paper. A portion is shown here for guidance regarding its form and content. Cols (1)-(3): identification, designation, and other identification, respectively. Cols (4)-(5): α and δ (J2000) coordinates are based on the H1 detection. Col (6): $v_{lsr}$ - the detected central frequency of the source. Col (7): $V_{rot}$ is redshift, defined as $z = (v_{rot} - v_{lsr})/v_{lsr}$, where $v_{lsr}$ is H I rest frequency at 1420.405751 MHz and $v_{rot}$ is observed frequency. Col (8): $V_{rot}$ = $cz$ - velocity in the optical reference frame. Cols (9)-(11): $V_{rot}$, $D_H$, and $D_L$ - recession velocity, Hubble distance, and luminosity distance, respectively (see Section 5.1). Col (12): $N_{HI,sen}$ - 1σ H1 column density sensitivity (see Section 3.2). Col (13): $S_{int}$ = $\sum S \times \Delta v$ - integrated flux (see Section 3.1). Col (14): $\log M_{HI}/M_\odot$ - observed H1 mass in logarithmic scale (see Section 5.2.2).

*aPartial H1 detection.

†Tidally interacting pair.

### Table 5. 2MASS K$_s$-band photometry, physical properties of sources, and relevant derived parameters.

<table>
<thead>
<tr>
<th>ID</th>
<th>Designation</th>
<th>$E(B-V)$ (mag)</th>
<th>$K_s$ (mag)</th>
<th>$err_{K_s}$ (mag)</th>
<th>$F_{K_s}$ (Jy)</th>
<th>$\nu L_\nu$ (× $10^9$ L$_\odot$)</th>
<th>$log M_{*}/M_\odot$ (M$_\odot$)</th>
<th>Morphology</th>
<th>$d_{25}$ (kpc)</th>
<th>DEF$_{HI}$</th>
<th>$i$ (°)</th>
<th>$w_{20}$ (km s$^{-1}$)</th>
<th>$V_{rot}$ (km s$^{-1}$)</th>
<th>Note</th>
</tr>
</thead>
<tbody>
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<td>1</td>
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<td>0.0561</td>
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<td>0.034</td>
<td>0.038</td>
<td>6.69</td>
<td>10.2</td>
<td>Sa</td>
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<td>69</td>
<td>57.6</td>
<td>–</td>
<td></td>
</tr>
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<td>0.0552</td>
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<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>44.4</td>
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<td></td>
</tr>
<tr>
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<td>–</td>
<td>–</td>
<td>–</td>
<td>dwarf?</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>85.0</td>
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Note: This table is available in its entirety as Supporting Information with the electronic version of the paper. A portion is shown here for guidance regarding its form and content. Cols (1) and (2): Identification and designation, respectively. Col (3): $E(B-V)$ - Galactic dust extinction (Schlafly & Finkbeiner 2011, see Section 5.2.1). Cols (4)-(6): $K_s$, $err_{K_s}$, and $F_{K_s}$ - intrinsic magnitude, error, and intrinsic flux density in 2MASS $K_s$ band, respectively (see Section 5.2.1). Col (7): $\nu L_\nu$ - luminosity in 2MASS $K_s$ band. Col (8): $log M_{*}/M_\odot$ - derived stellar mass in logarithmic scale (see Section 5.2.1). Col (9): Morphology - retrieved from NED or SIMBAD. Col (10): Calculated B-band optical isophotal diameter measured at 25 mag arcsec$^{-2}$ - $d_{25} = D_H \tan(B_{25,maj})$, where $B_{25,maj}$ is homogenized major axis from NED (see Section 5.2.2). Col (11): DEF$_{HI}$ - H1 deficiency (see Section 5.2.2). Col (12): $i$ - inclination as given by minor/major axial ratio to 3σ in 2MASS $K_s$ band. Cols (13)-(14): $w_{20}$ and $V_{rot}$ - corrected $w_{20}$ linewidth provided by SoFIA and calculated rotational velocity from the H1 linewidth, respectively (see Section 5.3).

*aPartial HI detection.

†Tidally interacting pair.
To verify the measured integrated fluxes from the ASKAP-12 data, we compare the spectra and integrated fluxes of a common (isolated) HI source between our study (WALLABY J132825–315124) and HIPASS (HIPASS 1328–31). For direct comparison, we extract the spectrum from the HIPASS sub-cube in the same way as we do for WALLABY J132825–315124. As shown in Fig. 7, the integrated fluxes in this case are consistent with each other.

We also compare the fluxes from ASKAP-12 with HIDEEP measurements in Fig. 8, where two outliers (labelled 1 and 2) are noted. We find that ASKAP source, IC 4290 (point 1), falls within the same beam as a single HIDEEP source (HIDEEP J1335–2801). Within the two overlapping beams of HIDEEP J1335–3025 and HIDEEP J1336–3030, there are three resolved ASKAP sources, namely ESO 444–77, 2MFGC 10975, and LEDA 715863. The value of $S_{\text{int,ASKAP}}$ for 2MFGC 10975 (1.1 ± 0.1 Jy km s$^{-1}$) is consistent with $S_{\text{int,HIDEEP}}$ (1.8 ± 0.2 Jy km s$^{-1}$). LEDA 715863 is not resolved by the HIDEEP survey and the flux contribution to the total $S_{\text{int,ASKAP}}$ is relatively small (0.6 ± 0.1 Jy km s$^{-1}$). The main deviation comes from the measured integrated fluxes of ESO 444–77. Examining the ASKAP-12 and HIDEEP cubes, we find that these sources are located in an area where S/N is low and only data from footprint B are present in the mosaicked cube. As mentioned in Section 2.3, the effect of solar interference is prominent in these beams, which may have resulted in additional flux. We should also point out that spectra of HIDEEP J1335–2801 and HIDEEP J1335–3025 reveal obvious baseline ripple, which can also affect the accuracy of integrated fluxes.

### 3.2 HI intensity and velocity field maps

We convert the HI intensity maps in units of Jy Hz from SOFIA to HI column density ($N_{\text{HI}}$) using equation 76 in M17:

$$\left( \frac{N_{\text{HI}}}{\text{cm}^{-2}} \right) = 2.33 \times 10^{20} (1 + z)^{3/2} S_{\text{HI}} \left( \frac{ab}{\text{arcsec}^2} \right)^{-1},$$

(1)

where $N_{\text{HI}}$ is the HI column density and $a$ and $b$ are beam angular major and minor axes, respectively. To better overlay a common set of contour levels on to optical images, we use HI column density sensitivity ($N_{\text{HI,sen}}$) levels. In Fig. 9, we show the ASKAP $N_{\text{HI,sen}}$ contours overlaid on to an optical $r$-band Digital Sky Survey 2 (DSS2r) or a combined $g$ and $r$-band Dark Energy Camera (DECam) image for all detected sources. DECam images are at higher resolution than DSS2r, hence preferred wherever available. The contour levels are 3, 5, 7, 9, 11, and 13 times $N_{\text{HI,sen}}$, where $N_{\text{HI,sen}}$ is given by

$$\left( \frac{S_{\text{HI,sen}}}{\text{Jy Hz}} \right) = \sigma_{\text{rms}} \times \sqrt{N_{\text{chan}} \times \Delta v},$$

(2)
Figure 9. Examples of integrated H I column density maps of individual sources overlaid on to DSS2r or DECam (Müller, Jerjen & Binggeli 2015) images. The contour levels are 3, 5, 7, 9, 11, and 13 $\times N_{\text{HI,sen}}$ in cm$^{-2}$, (refer Table 4), which corresponds to colour scale of blue, orange, red, purple, pink, and olive green, respectively. The synthesized beam of 45 arcsec $\times$ 35 arcsec is plotted at the bottom left corner of each sub-plot as a reference. Note. This figure is published its entirety as Supporting Information with the electronic version of the paper. A portion is shown here.

Figure 10. Examples of velocity field maps of individual sources, where only pixels above $3 \times N_{\text{HI,sen}}$ are plotted. The synthesized beam of 45 arcsec $\times$ 35 arcsec is plotted at the bottom left corner of each sub-plot. Note. This figure is published its entirety as Supporting Information with the electronic version of the paper. A portion is shown here.

where $\sigma_{\text{RMS}}$ is the local RMS in Jy, $N_{\text{chan}}$ is the total number of channels that contain H I emission and $\Delta$$\nu$ is the channel width of ASKAP data. Then, the conversion from flux density (sensitivity) to $N_{\text{HI}}$ (sensitivity) is given in equation (1). This gives an upper limit for $N_{\text{HI,sen}}$ because not all pixels contain the same number of channels with H I emission.

In Fig. 10, we present the 1st moment velocity field maps of all sources. Pixels below $3 \times N_{\text{HI,sen}}$ are masked out. We note that most of the detected galaxies are marginally resolved.

4 GROUP MEMBERSHIP

Galaxy group membership depends strongly on the selection criteria and algorithms. While the G93 catalogue is constructed with a combination of methods, it is limited to magnitude ($B_{\text{mag}} < 14.0$) and only to the nearby Universe ($V_{\text{CMB}} \leq 5500$ km s$^{-1}$). More recent studies of galaxy group membership have been carried out by Yang et al. (2007), who used SDSS (York et al. 2000) and 2dFGRS (Colless et al. 2001) data, with vastly more spectroscopic redshifts compared with G93. Tully et al. (2013) has also constructed a galaxy group catalogue based on a compilation of distances and peculiar velocities of $>8000$ galaxies. Both Yang’s and Tully’s catalogues are constructed with different selection criteria and scientific goals, and are not directly comparable.

Given that the H I detected galaxies in our study are close in projected separation and in velocity to the existing galaxies defining LGG 351 (see Fig. 5), they are likely to all be members of an extended group. To determine if associated with LGG 351, we calculate the velocity difference between each galaxy and the mean central velocity of LGG 351 ($\Delta V = V_{\text{opt}} - \bar{V}$) and then compare it to $3\sigma_{\text{vel}}$. With this criterion, we exclude IC 4290 from being part of the group. Re-iterating with the new $\bar{V}$ and $\sigma_{\text{vel}}$, we do not find any other exclusions. Many of the detected H I sources have also been identified as part of the TSK group 2543 (Tully et al. 2008, 2013). This part of the local Universe is markedly overdense, due to being within the extended Hydra-Centaurus supercluster.

4.1 Dwarf galaxies

Dwarf galaxies play an important role as they are the building blocks of galaxies. In the context of H I, dwarf galaxies in the Local Group as well as the newly discovered satellites of the Milky Way and the Andromeda galaxy are H I deficient (Greffich & Putman 2009; Westmeier et al. 2015). This is in part due to tidal and ram-pressure effects in the vicinity of massive galaxies in the Local Group. The discovery of gas-rich ultra-faint dwarf galaxies, such as Leo P (Giovanelli et al. 2013; McQuinn et al. 2015) suggests that the census of dwarf satellites remains incomplete due to sensitivity limitations of instruments at both in optical and radio wavelengths.

A search for ultra-faint dwarf galaxies with DECam down to $M_v \sim -10$ mag has been conducted in the vicinity of the Centaurus group, which consists of the Centaurus A and M 83 subgroups (Müller...
we scale the $M_\odot$ and is at the low end of HI masses $\lesssim 10^8 \, M_\odot$. Stellar masses of galaxies in our study are of the order of $10^{8.5-11} \, M_\odot$.

5 PHYSICAL PARAMETERS

5.1 Luminosity and Hubble distances

We calculate the Hubble distance ($D_H$), which is defined as $D_H = V_{\text{CMB}}/H_0$, where

$$V_{\text{CMB}} = V_{\text{opt}} + V_{\text{apex}} \left[ \sin b \sin b_{\text{apex}} + \cos b \cos b_{\text{apex}} \cos(l - l_{\text{apex}}) \right]$$

is the recession velocity in the cosmic microwave background (CMB) reference frame, $l_{\text{apex}} = 264.14^\circ$, $b_{\text{apex}} = 48.26^\circ$, and $V_{\text{apex}} = 371.0 \, \text{km s}^{-1}$ (Fixsen et al. 1996). Galactic coordinates of galaxies are represented by $l$ and $b$. We can also calculate the luminosity distance, $D_L = (1 + z)D_H$, where $D_L$ is the co-moving distance. At low redshift, $D_L \approx D_H$. In Fig. 11, we show the $V_{\text{CMB}}$ distribution of H I detected galaxies in the field. The distribution peaks at $\sim 4600 \, \text{km s}^{-1}$ and the mean LGG 351 $V_{\text{CMB}}$ is $4500 \, \text{km s}^{-1}$, which corresponds to a mean $D_L$ of 63.2 Mpc.

5.2 Masses

5.2.1 Stellar mass

Stellar masses can be derived by using colour–mass relations. Near-infrared (NIR) luminosities of galaxies are useful for estimating the mass-to-light ratio (M/L) as they only vary slightly across a wide range of star formation histories (Bell & de Jong 2001; Bell et al. 2003). To make use of the existing NIR colour–stellar mass relations, we first obtain the $K_s$ colours of our galaxies from the 2MASS extended source catalogue (Skrutskie et al. 2006). Using the wavelength-dependent extinction law of Cardelli, Clayton & Mathis (1989), we derive $A_{K_s} = 0.36 \, E(B - V)$. The Galactic dust extinction, $E(B - V)$, of each galaxy is estimated using the re-calibrated SFD all-sky extinction maps (Schlegel, Finkbeiner & Davis 1998; Schlafly & Finkbeiner 2011). Following the procedure as described in section 6 of For, Koribalski & Jarrett (2012), we obtain the intrinsic fluxes in the $K_s$ band by adopting the zero-magnitude flux of the star Vega ($F_0$). The corresponding $F_0$ is 666.7 ± 12.6 Jy (Cohen, Wheaton & Megeath 2003). We also estimate the $k$-correction for the $K_s$ band (Chilingarian, Melchior & Zolotukhin 2010; Chilingarian & Zolotukhin 2012), which ranges from $-0.03$ to $-0.05$ mag. The correction is smaller than a typical $K_s$ mag error of 0.09 and hence is considered to be negligible. Subsequently, we obtain the stellar masses by employing the 2MASS $K_s$ band to stellar mass relation as given in equation 3 of Wen et al. (2013):

$$\log \left( \frac{M_*}{M_\odot} \right) = (-0.498 \pm 0.002) + (1.105 \pm 0.001) \log \left( \frac{vL_\nu}{L_\odot} \right).$$

This relation uses the Salpeter (1955) IMF. For comparing with the literature, where the Chabrier (2003) IMF is commonly used, we scale the $M_*$ values by a factor of 0.61 (Madau & Dickinson 2014). Stellar masses of galaxies in our study are of the order of $10^{8.5-11} \, M_\odot$. 

Figure 11. Histogram of recession velocities ($V_{\text{CMB}}$). Orange and blue represent galaxies in G93 LGG 351 catalogue and the newly identified galaxies, respectively. IC 4290, which is probably not a group member is represented in green.

Figure 12. Histogram of H I masses. Orange and blue represent galaxies in G93 LGG 351 catalogue and the newly identified galaxies, respectively. Two peaks are seen in this distribution, with one at the low-mass end and the other at the high-mass end. IC 4290 is represented in green.

et al. 2015; Müller, Jerjen & Binggeli 2017). They find 41 new dwarf galaxy candidates, which nearly doubles the number of known galaxies in the Centaurus group. Confirmation of their membership relies on the tip of the red giant branch distance measurements and the surface brightness fluctuation method (Tonry & Schneider 1988), but only a handful of them have been measured to date (Müller, Rejkuba & Jerjen 2018).

The H I spectral line provides another way to estimate distances (see Section 5.1) if these ultra-faint dwarf galaxies contain H I. Analysis of HIPASS data shows no H I detection and gives an upper limit of $M_{\text{HI}} < 8.5 \times 10^6 \, M_\odot$ to these dwarf galaxy candidates (Müller et al. 2017). Nevertheless, we search for an H I signature of some of these ultra-faint dwarf galaxies given that they fall within the studied area. We find a detection of H I for dw 1328–29 (#11), which has $M_{\text{HI}}$ of $3.4 \times 10^6 \, M_\odot$ and is at the low end of H I masses of galaxies within LGG 351 (see Fig. 12). This also rules out dw 1328–29 as a member of the M 83 group. We also identify LEDA 98614 (#14) as a dwarf candidate based on the combined $g$ and $r$-band DECam image.
projected separation to less massive galaxies. The majority of the most massive galaxies are in close proximity when DEF$_{HI} \leq -0.5$. These correspond to 70 per cent less or more massive galaxies as compared to isolated galaxies of the same morphological type and diameter.

We obtain the morphological type and calculate the $D_{maj}$ based on the $B_{maj}$ semimajor axis ($B_{maj}$) from NED. There are only 11 galaxies in our study with available morphological classifications, $B_{maj}$ measurements and HI masses. Among them, NGC 5152 (DEF$_{HI} = 0.98$) and ESO 444–G047 (DEF$_{HI} = -0.56$) are considered H I deficient and H I excess. The finding is similar to the NGC 7162 group, in which majority of the group members are H I normal with one galaxy (ESO 288–G025) being H I deficient (DEF$_{HI} = 0.50$) (R197). However, we note that our sample is small for statistical interpretation.

A recent study on H I scaling relations derived from isolated galaxies, which may not be strongly affected by the environment, is described in Jones et al. (2018). The samples are drawn from the Analysis of the interstellar Medium of Isolated GaLaxies project (AMIGA; Verdes-Montenegro et al. 2005) and are divided into three morphological types (early-type, late-type, and very late-type galaxies). The derived coefficient ($b$ as in equation 7) for the late-type galaxies in this study is larger than in Solanes et al. (1996) but is more consistent with the study of Haynes & Giovanelli (1984). The difference may indicate an influence on environment. We refer the readers to Jones et al. (2018) for a detailed comparison and discussion of different H I scaling relations and note that the derived scaling relations are subject to the selection effects.

5.3 Rotational velocity

The rotational velocity of a galaxy can be calculated from the H I linewidth, as given by

$$V_{rot} (\text{km s}^{-1}) = \frac{w_{20}^{2}}{2 \sin i},$$

where $w_{20} = \sqrt{w_{20}^{2} - (\Delta s/2)}$ is the H I linewidth measured at 20 per cent of the level of each peak, and corrected for instrumental broadening ($\Delta s$) and $i$ is the inclination. According to Springob et al. (2005), $\Delta s = 2 \Delta v$, where $\Delta v$ is assumed to be equal to, or close to, the channel width in velocity space, and $\lambda$ is the broadening parameter. We adopt $\lambda = 0.5$, which is appropriate for typical profile shapes. We can compute the inclination using the following equation:

$$\cos i = \sqrt{\frac{(b_{maj}/b_{maj})^{2} - q^{2}}{1 - q^{2}}}.$$

where $b_{maj}/b_{maj}$ is the axial ratio of minor-to-major axes from 2MASS $K_{S}$ band and $q$ is the intrinsic axial ratio. We adopt a constant $q$ of 0.2 (Tully et al. 2009) because the difference in $q$ values is small between different morphological types (see discussion in Catinella et al. 2012; hereafter C12). The calculated $V_{rot}$ are given in Table 5.

6 COMPARISON WITH SCALING RELATIONS

A series of studies on H I scaling relations have been carried out as part of the GALEX Arecibo SDSS Survey (GASS; see Catinella et al. 2013 and references therein) and extended GASS (xGASS; see Catinella et al. 2018 and references therein). GASS

\footnote{For consistency, we revise their calculation and find that those galaxies previously reported to be H I excess are H I normal.}
is a large Arecibo H I survey that targets galaxies with stellar masses greater than $10^{10} \, M_\odot$ and redshifts 0.025 < z < 0.05. xGASS is an extension of GASS to include lower stellar mass galaxies ($10^9 < M_* < 10^{11.5} \, M_\odot$) and redshifts 0.01 < z < 0.05. In this section, we compare our derived parameters to H I scaling relations.

6.1 Atomic gas fraction scaling relation

The atomic gas fraction ($M_{HI}/M_*$) varies as a function of integrated galaxy properties and decreases with $M_*$. Massive galaxies ($M_* > 10^{10} \, M_\odot$) within large haloes ($M_h = 10^{13}-10^{14} \, M_\odot$) have less H I gas fractions than those galaxies with similar $M_*$ in smaller haloes (Catinella et al. 2015). In Fig. 14, we overplot our data (red triangles) on to the atomic gas fraction scaling relation that is based on the GASS and xGASS data (grey circles). The average atomic gas fraction of isolated central galaxies (blue diamonds) and central galaxies in group (green squares) from the study of Janowiecki et al. (2017) is also shown. They find that the average gas fraction in central galaxies in group below $10^{10.2} \, M_\odot$ is $\sim$0.3 dex higher than isolated central galaxies. The majority of our galaxies ($M_* > 10^9 \, M_\odot$) are above the green line, which consistent with the Janowiecki’s result. It is unclear if the discrepancy of gas fraction between isolated central galaxies and central galaxies in group remains below $10^9 \, M_\odot$. Further investigation is needed.

However, examining the atomic gas fraction scaling relation that is derived from the Herschel Reference Survey (HRS; Boselli et al. 2010), which covers stellar masses below $10^9 \, M_\odot$ (see fig. 1 in Cortese et al. 2011), we find that the low-mass galaxies in this study follow the same general trend. The two low stellar-mass galaxies, WALLABY J133035−293421 and 2MFGC 10975 with log $M_*/M_\odot$ > 0, indicate that gas is inefficiently converting into stars.

6.2 Baryonic Tully–Fisher relation

The baryonic mass is the sum of gas and stellar masses, in which the total gas mass is mostly composed of molecular and neutral hydrogen. The contribution from H I is small for the range of stellar masses being studied in C12.

As such, we use $M_{gas} = 1.4 \, M_{HI}$, where a 1.4 correction factor is applied to account for helium and metals. In Fig. 15, we overplot our data (filled red triangles) on to the GASS and xGASS data (grey circles) for the Baryonic Tully–Fisher (BTF) relation. The dashed line represents the fit from C12. We find that $V_{rot}$ for five galaxies (2MFGC 10975, NGC 5152, IC 4290, NGC 5150, and NGC 5182) deviates by more than 0.1 dex from the fitted BTF scaling relation. However, we note that the scatter for the C12 samples is 0.127 dex, with a smaller scatter of 0.076 dex for their subset samples. The scatter here is comparable with the main uncertainties contributed by observational errors in $w_{20}$ and inclination.

7 STAR FORMATION RATE

Ultraviolet (UV) luminosity is an excellent tracer of recent star formation (<100 Myr) but it is affected by dust attenuation. Most UV continuum photons are absorbed by dust and then re-emitted in the mid-infrared (MIR) wavelength. In this study, we use the Galaxy Evolution Explorer (GALEX; Martin et al. 2005; Morrissey et al. 2007) and Widefield Infrared Survey Explorer (WISE; Wright et al. 2010) data to estimate the dust un-attenuated and attenuated part of the total SFRs of our galaxies, respectively. It is a known issue that the fluxes obtained from the GALEX pipeline are not accurate for extended objects. This is due to the use of small Kron-like elliptical apertures that results in extended sources being broken up into pieces. To circumvent this issue, Seibert et al. (2012) created the GALEX Source Catalog in the Kepler Field (GCAT7), in which specific apertures were used to obtain fluxes for the identified extended sources. We obtain the GALEX NUV AB magnitudes by matching our galaxies’ coordinates to the GCAT via the GALEX CasJobs interface. GCAT provides an homogeneous photometry, which includes data from the GALEX Medium Imaging Survey (MIS), the All-Sky Imaging Survey (AIS), and the Deep Imaging Survey (DIS) through GR7. We adopt the MIS over AIS magnitudes whenever available. The associated magnitude uncertainty is $\sim$0.15 mag (Seibert et al. 2012).

We follow the same procedure as described in Section 5.2.1 to correct NUV AB magnitudes for the Galactic dust extinction, where $A_{NUV} = 8.2 \times (B−V)$ (Wyder et al. 2007) is adopted. The intrinsic monochromatic flux $F_{\nu, mono}$ in units of erg s$^{-1}$ cm$^{-2}$ Hz$^{-1}$

7https://archive.stsci.edu/prepds/gcat/
We convert the intrinsic NUV fluxes to luminosities. The measurement.

To investigate if our galaxies are actively forming stars, we
compare SFRs with the fitted SFR–$M_*$ relations derived from a sample of low surface brightness galaxies (LSBG; McGaugh, Schombert & Lelli 2017) and star-forming galaxies in xCOLD GASS (extended CO Legacy Database for GASS; Saintonge et al. 2017). This relation forms the star formation main sequence (SFMS), where galaxies are undergoing exceptionally high star formation if lie above the SFMS and are quenched if lie below the SFMS. We plot this relation in Fig. 16. We overplot our data (filled orange and green squares), the LSBG data (blue dots) and the xCOLD GASS data (grey and black dots) on to the SFR–$M_*$ plane. The filled orange and green squares represent the total SFR and the lower limit, respectively. Errors in total SFR are relatively small compared with the size of the symbol. The grey and black dots represent the xCOLD GASS galaxies and those without CO(1–0) detections, respectively. The study of McGaugh et al. (2017) covers a stellar mass range of $10^7$–$10^{10}$ $M_\odot$ and extends the SFR down to $10^{-4}$ $M_\odot$ yr$^{-1}$. They find a steep slope (solid and dark dashed line) for the SFR–$M_*$ relation but a ‘flatter’ slope if compared to the relation derived from the Spitzer Infrared Nearby Galaxies Survey sample (not shown), which mostly consists of massive spirals ($M_*$ $\sim 10^{10} M_\odot$) (Kennicutt et al. 2009). The derived SFR–$M_*$ relation (red line) using a large sample of COLD GASS shows a similar flattening (Saintonge et al. 2016). They find that the flattening of SFMS is due to the global decrease of cold gas reservoirs in galaxies rather than the inefficiency in converting cold atomic gas to stars. We find NGC 5152 and IC 4290 (labelled as 2 and 3, respectively in Fig. 16) fall below the SFMS relation suggesting quenching in star formation. In fact, NGC 5152 is deficient in H I. While LEDA 47428 (labelled as 1) is positioned below the SFMS of both McGaugh and Saintonge’s relations, we cannot conclude if it is also quenched in star formation. This is because its SFR is a lower limit, but within the scatter of McGaugh’s SFMS relation. NGC 5135 (labelled 4), is a known starburst galaxy, which its SFR is expected to be above the SFMS. Table 6 gives a summary of the photometry and the calculated SFRs of the 34 galaxies.
Table 6. NUV and MIR photometries of sources and associated star formation rate.

<table>
<thead>
<tr>
<th>Designation</th>
<th>NUV$_{\text{MIS}}$ (AB mag)</th>
<th>NUV$_{\text{AIS}}$ (AB mag)</th>
<th>W4 (mag)</th>
<th>err$_{W4}$ (mag)</th>
<th>SFR$<em>{\text{NUV}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
<th>SFR$<em>{\text{W4}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
<th>SFR (M$_{\odot}$ yr$^{-1}$)</th>
<th>err$<em>{\text{SFR}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>WALLABY J132514—303304</td>
<td>–</td>
<td>–</td>
<td>4.80</td>
<td>0.06</td>
<td>–</td>
<td>1.27</td>
<td>&gt;1.27</td>
<td>–</td>
</tr>
<tr>
<td>WALLABY J132533—300521</td>
<td>17.98</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>0.08</td>
<td>–</td>
<td>0.08</td>
<td>0.02</td>
</tr>
<tr>
<td>WALLABY J132541—295013</td>
<td>–</td>
<td>1.55</td>
<td>0.06</td>
<td>–</td>
<td>25.37</td>
<td>&gt;25.37</td>
<td>–</td>
<td>0.02</td>
</tr>
<tr>
<td>WALLABY J132627—301909</td>
<td>18.18</td>
<td>–</td>
<td>–</td>
<td>0.07</td>
<td>–</td>
<td>0.07</td>
<td>0.07</td>
<td>0.02</td>
</tr>
<tr>
<td>WALLABY J132644—300031</td>
<td>17.90</td>
<td>17.90</td>
<td>–</td>
<td>–</td>
<td>0.07</td>
<td>–</td>
<td>0.07</td>
<td>0.02</td>
</tr>
</tbody>
</table>

Note. This table is available in its entirety as Supporting Information with the electronic version of the paper. A portion is shown here for guidance regarding its form and content. Col (1): Designation. Cols (2)–(3): NUV$_{\text{MIS}}$, NUV$_{\text{AIS}}$ – GALEX NUV AB magnitude obtained from the Medium Imaging Survey and All-sky Imaging Survey. Value has been corrected for the Galactic dust extinction. Cols (4)–(5): WISE band-4 22 μm (mid-IR) magnitude and error. Value has been corrected for Galactic dust extinction. Cols (6)–(7): Calculated star formation rates based on the NUV and mid-IR fluxes, respectively. Cols (8)–(9): Total star formation rate = SFR$_{\text{NUV}}$ + SFR$_{\text{W4}}$ and its error, respectively.

$^*$Tidally interacting pair.

Figure 17. Rotation curves derived from 2DBAT. Circles, crosses, and squares represent receding, approaching, and both sides of the rotation curves. The dashed lines show the fit to the ‘flat’ part of the rotation curves used for estimating the errors.

To determine the rotation velocities of these four galaxies, we take the average values of data points that are on the ‘flat’ part of the rotation curves. We perform a linear least-squares fit to these data points and estimate uncertainties. The red dashed lines in Fig. 17 indicate the data points that are used for fitting and averaging. A summary of the derived parameters is given in Table 7. The $V_{\text{rot}}$ of IC 4275 has a large uncertainty (about ±160 km s$^{-1}$) due to its low derived inclination, which is different from the optical inclination (∼60$^\circ$). In Fig. 18, we show the input single Gaussian extracted velocity field maps (left-hand panels), tilted-ring models (middle panels) and residual velocity fields (right-hand panels). These four galaxies exhibit a typical velocity field that is dominated by circular kinematic structures of our galaxies.

$^8$PA is defined as the angle between the major axis of the receding half of the galaxy and north.
of slight distortion. Extended H I gas from the north-east corner of all group members in LGG 351 of G93 catalogue shows evidence In contrast, ESO 444−G047 and velocity field map of IC 4275 suggests that this galaxy is warped. (see Table 7) are comparable with the tilted-ring models of IC 4248, IC 4275, and ESO 444−G047 −1 V cannot be made for LEDA 98836 as there is no optical inclination and H I linewidth after taking into account the uncertainties (see Section 5.3 and Table 5). A comparison with only footprint A coverage and the exclusion of beam 31 of footprint B.

There are many H I galaxies around IC 4275 (#28) with LEDA 3082253 (#30) and WALLABY J133201−294119 (#29) being closest to IC 4275. LEDA 3082253 shows an H I extension towards the south-east corner. The optical DECam image also reveals an asymmetry in the stellar distribution, with faint structures extending in the same direction as the H I. Due to their close proximity, tidal interaction is probably the cause of the H I distortion as seen in their extended H I disc. There is no H I debris found. WALLABY J133201−294119 is a new galaxy discovered serendipitously through the deep DECam image.

NGC 5152 (spiral) and NGC 5153 (elliptical) are an interacting pair (Arp & Madore 1987). This is particularly interesting as substantial H I debris are often found near massive, early-type galaxies (see e.g. Chung et al. 2009; Saponara et al. 2018). The pair is also part of the subgroup containing NGC 5150. In Fig. 19 (left-hand panel), we show the H I contours from the H I survey overlaid on to the DSS2 image, where the H I detection is unresolved and identified as H I survey footprint B.

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debris is unexpected, but it is likely that the diffuse HI emission has a lower HI column density ($<10^{20}$ cm$^{-2}$) that cannot be detected by WALLABY. Another interesting fact for this subgroup is that two tidal dwarf candidates (d1 and d2) have also been reported by Weilbacher et al. (2000) (see bottom panel). The formation of these tidal dwarf galaxies could only be possible via galaxy interactions, when a large amount of gas is expelled and re-collapsed to form such systems.

IC 4248 (#6) is a Sab galaxy. Its optical image shows an interesting ram-pressure/tidally stripped morphology with a few visible ‘tentacles’. The HI morphology also shows a distorted structure with a asymmetric morphology similar to the optical morphology. It is known to be interacting with NGC 5135 (#3) (de Vaucouleurs, de Vaucouleurs & Corwin 1976). Unfortunately, we do not find any conclusive evidence of ram-pressure stripping from the tilted-ring fitting given that there is no change in PA. Further investigation is possible with imaging that resolves the inner and outer disc and further down to $10^{19}$ cm$^{-2}$.

While we only see hints of tidal interaction within LGG 351 based on HI morphologies, a new interacting pair in the vicinity, WALLABY J133002–272832 (#17), is discovered. This pair includes LEDA 47428 and ESO 509–G032 (an edge-on galaxy). The DECam image shows a spiral structure at the far side of ESO 509–G032 and a faint extended thin stellar disc towards LEDA 47428. The galaxy pair is located at the northern end of our observed field.

10 FURTHER DISCUSSION

Galaxies in denser cluster environments tend to be deficient in HI as compared to field galaxies (Solanes et al. 2001). Studies of the Fornax (Waugh et al. 2002) and Coma clusters (Bravo-Alfaro et al. 2000) have provided evidence of a decrease of HI with distance from the cluster centre (see also Haynes & Giovanelli 1984). They also show that the optical discs of galaxies near the cluster centre extend beyond their HI discs. This strongly suggests that interactions with the hot IGM (i.e. ram-pressure stripping) is likely the dominant gas removal mechanism in the cluster environment.

Nevertheless, HI deficiency in galaxies is not unique to the cluster environment. It is also found in galaxies of compact and loose groups (see e.g. Verdes-Montenegro et al. 2001; Kilborn et al. 2005). Detection of loose X-ray emission associated with the IGM in Hickson compact groups suggests that tidal interaction rather than ram-pressure stripping is the dominant gas removal mechanism in compact groups (Rasmussen et al. 2008).
As for loose galaxy groups, asymmetries in the outer H I discs of NGC 300 and NGC 55 in the Sculptor group suggest the presence of ram-pressure stripping (Westmeier et al. 2011; Westmeier, Koribalski & Braun 2013). These findings imply that both ram-pressure stripping and tidal interaction play a significant role in removing gas in galaxy groups, but which is the dominant mechanism remains an open question.

There are three galaxies (NGC 5153 and IC 4251 and PGC 47224) in LGG 351 that do not have an H I detection in the observed field. NGC 5152 (in the subgroup of NGC 5150) is considered to be H I deficient and is likely as a result of gas being pulled out to form the tidal dwarfs (see Section 9). Given that IC 4251 and PGC 47224 are located relatively close to the NGC 5150 subgroup and evidence of tidal interaction within the subgroup is seen (see Section 9), we suggest that tidal interaction is likely the main gas removal mechanism for these galaxies despite the lack of H I debris. Further investigation of the hot gas in the region might rule out or support ram-pressure stripping mechanism for these galaxies.

**11 SUMMARY AND CONCLUSIONS**

WALLABY early science data have allowed us to test and to optimize the analysis techniques and observing strategy for WALLABY. In this process, we have improved continuum subtraction and have found optimal observing window for day time observation to minimize solar interference. We have also tested and demonstrated the performance of ASKAPsoft, increasing the capability of new versions to deal with early science and full WALLABY data. A statistical analysis of data quality, influence of solar interference and comparison with HIPASS and HIDEEP integrated fluxes is made. The assessment of the effect of solar interference could not be done in previous early science data because there were observed at night. Overall, ASKAP has recovered all the fluxes when compared to single-dish HIDEEP data. Two outliers will need further reassessment with full WALLABY data. We perform source finding using SOFIA and obtain a catalogue of 40 H I sources. Among these sources, two are new H I galaxies that do not have optical identifications and one is a tidally interacting pair. This is the largest WALLABY sample to date. All confirmed H I detections have optical counterparts. We use our measurements to refine LGG 351 group membership, adding 32 new members (including a tidally interacting pair). Further observations are likely to further increase group membership, and resolve membership issues in LGG 351 and the relationship with TSK group 2543.

Due to the angular resolution, we cannot study the small-scale kinematic structure. We also present the extracted velocity fields, model, and residual maps for these four galaxies. Investigating the H I morphology of the LGG 351 galaxies, we find evidence of tidal interaction in the NGC 5150 subgroup, NGC 5182, LEDA 3082253 and IC 4248. There is no H I debris found in our studied area. It is likely that any extended diffuse emission has been resolved out by ASKAP with the early science configuration. Full ASKAP-36 WALLABY observations or follow-up targeted H I observations with higher sensitivity would be useful in the search for extended diffuse emission and allow further study the galaxies in greater detail.

This study has shown the excellent performance of SOFIA in searching for sources in ASKAP H I data cubes and the feasibility of using 2DBAT to derive the rotation curves of marginally resolved galaxies. WALLABY early science has also proven to be sensitive enough to detect low-mass dwarf galaxies of the order of $10^8 M_\odot$ out to a distance of ~60 Mpc. Complementary multiwavelength data are shown to be important in order to study physical processes and their relations with the environment. Our sample is small for some of the derived parameters, such as the star formation and stellar mass, due to the lack of UV, infrared, and optical photometry. The ongoing Skymapper survey, VISTA Hemisphere Survey and extended source catalogue from WISE could potentially mitigate the issue of sample size.

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WALLABY early science – V. The LGG 351

Tully R. B. et al., 2013, AJ, 146, 86

SUPPORTING INFORMATION

Supplementary data are available at MNRAS online.

Figure 6. Spectrum of individual sources obtain from SOFIA.

Figure 9. Integrated H I column density maps of individual sources overlaid on to DSS2r or DECam images.

Figure 10. Velocity field maps of individual sources.

Figure 18. Input velocity field maps, tilted ring models, and residual velocity fields.

Table 4. Source catalogue and derived parameters.

Table 5. Photometry of sources and relevant derived parameters.

Table 6. NUV and MIR photometries of sources and associated star formation rate.

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